

The stellar background noise in intensity and velocity

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Tim Bedding

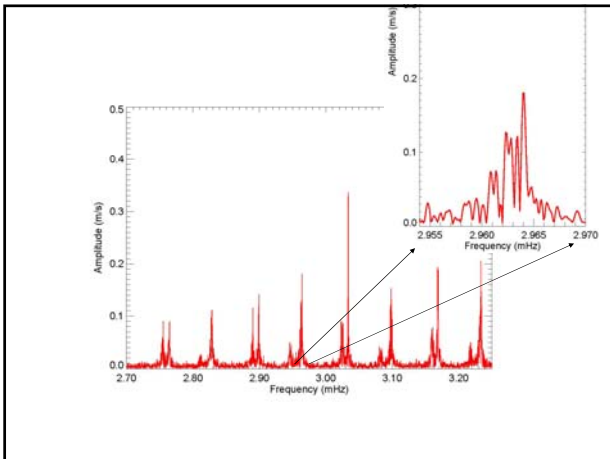
School of Physics, University of Sydney, Australia

Properties of the p-mode Oscillations

- Peak power: level and frequency
- p-mode frequencies
- Damping time (mode life)
- Noise background (granulation)

Detection of low-Mass planets

- Amplitude from Sun-Earth: 9 cm/s, 1 yr

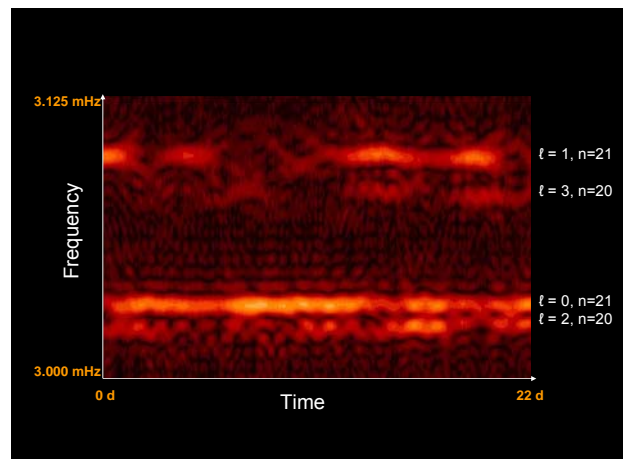
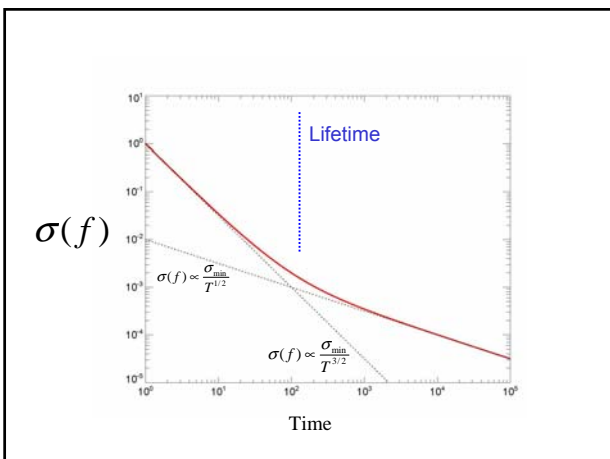


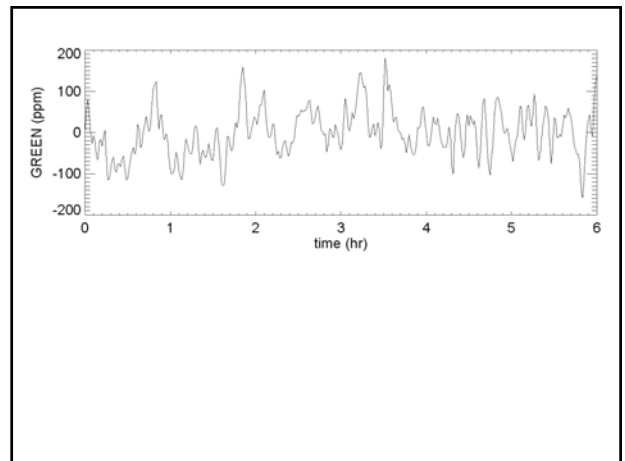
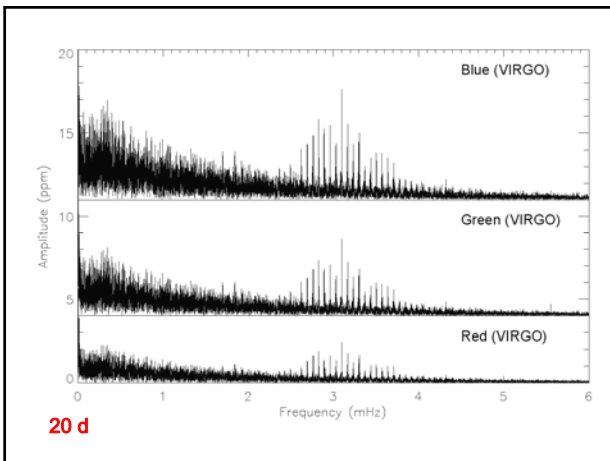
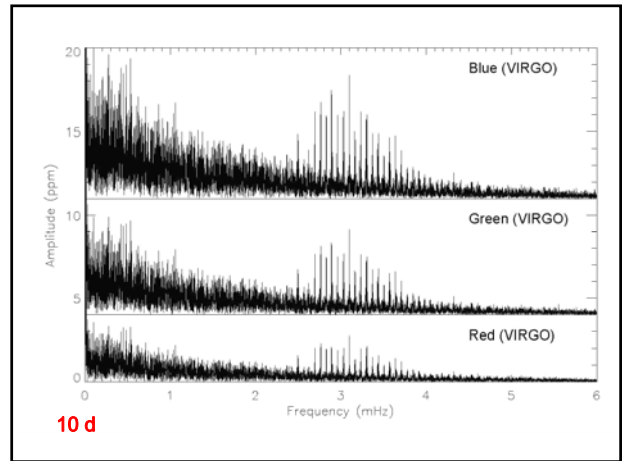
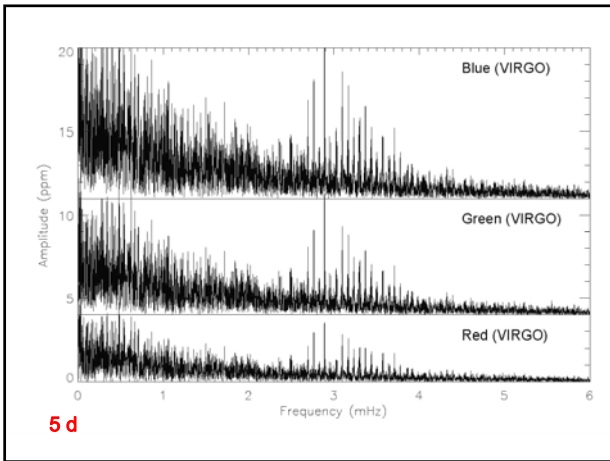
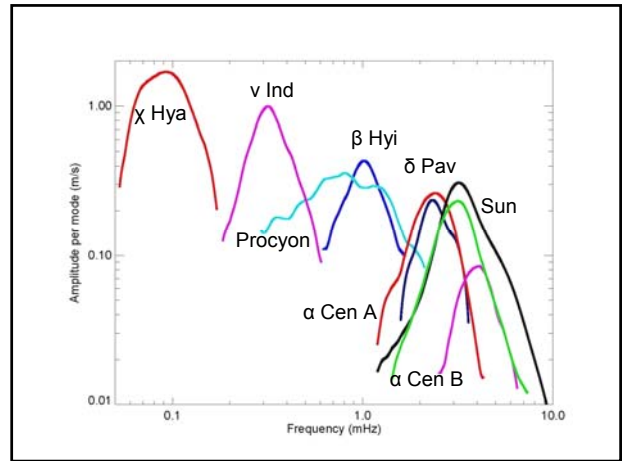
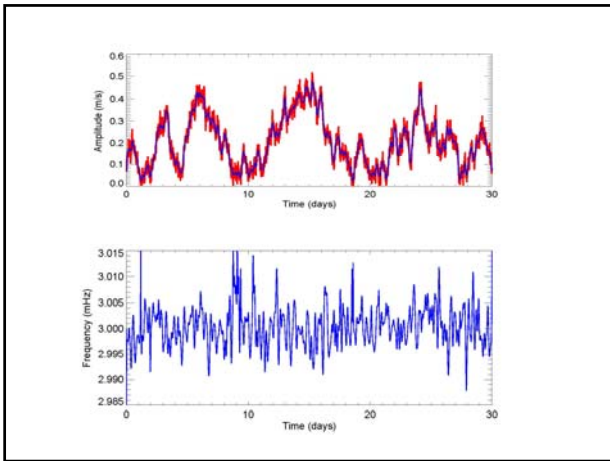
Coherent modes

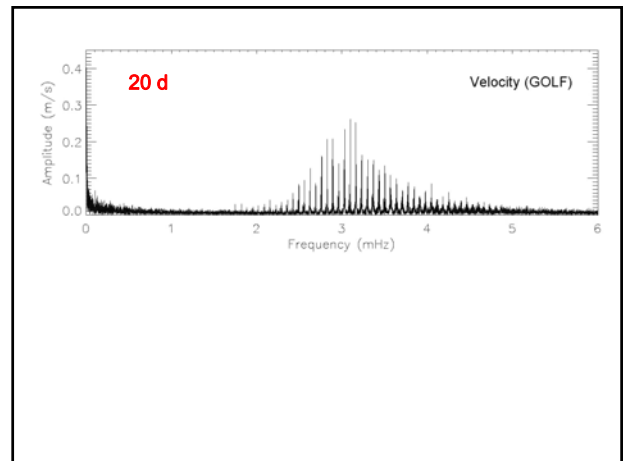
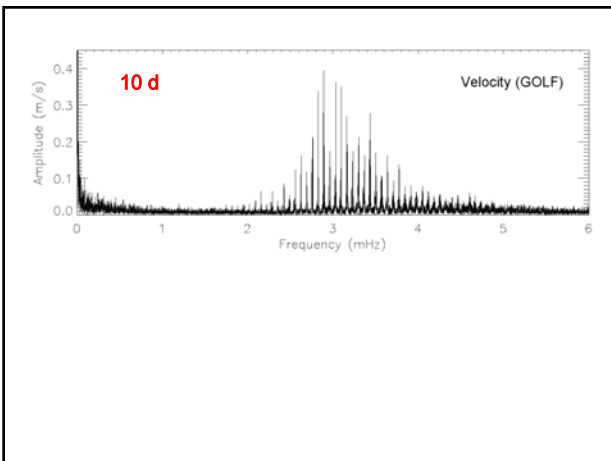
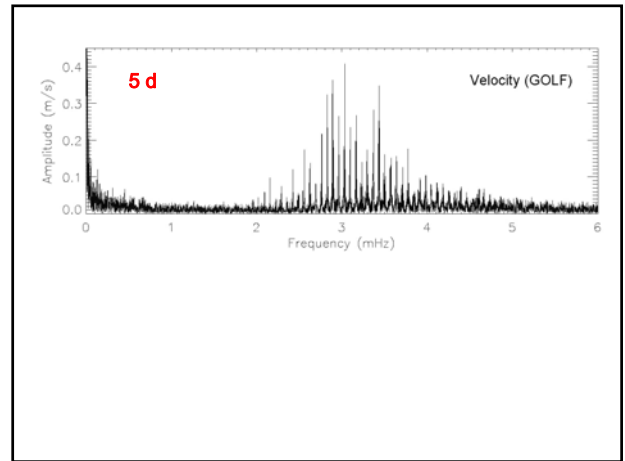
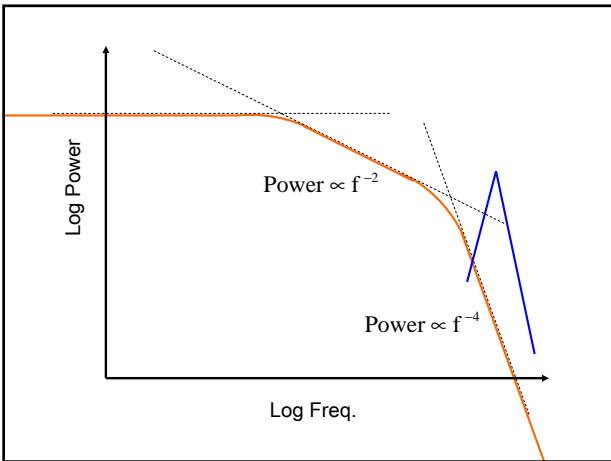
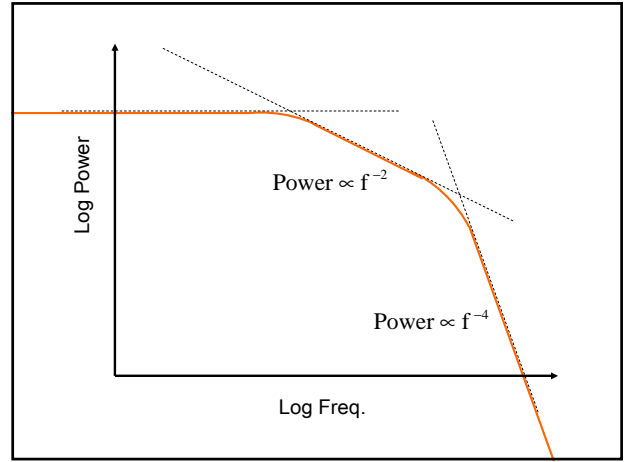
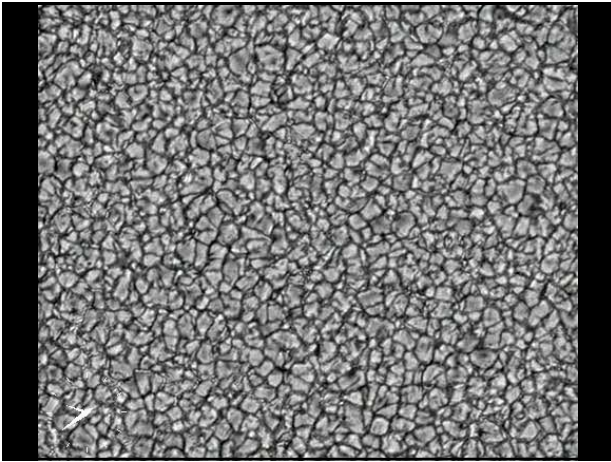
$$\sigma(f) \propto \frac{\sigma_{\min}}{T^{3/2}}$$

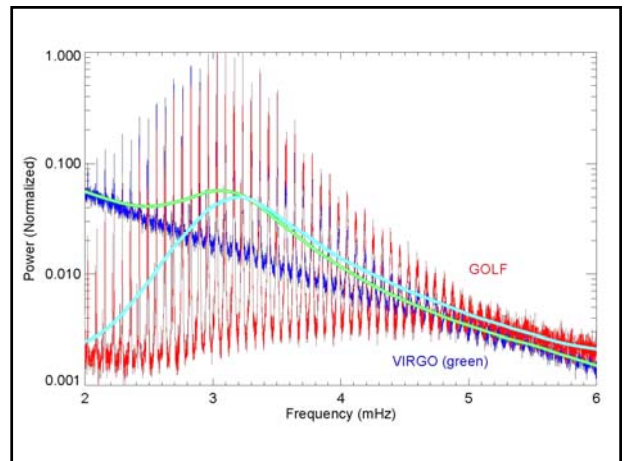
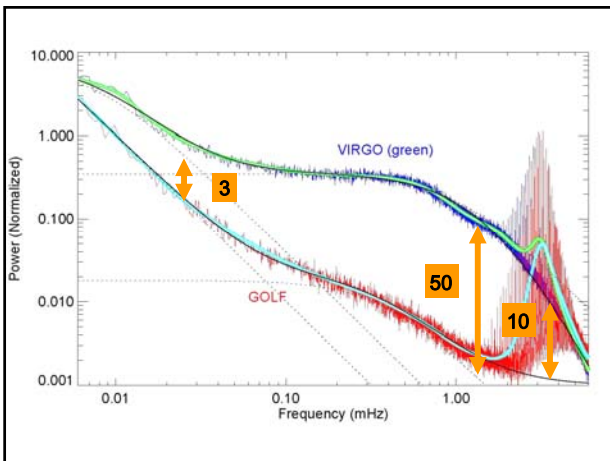
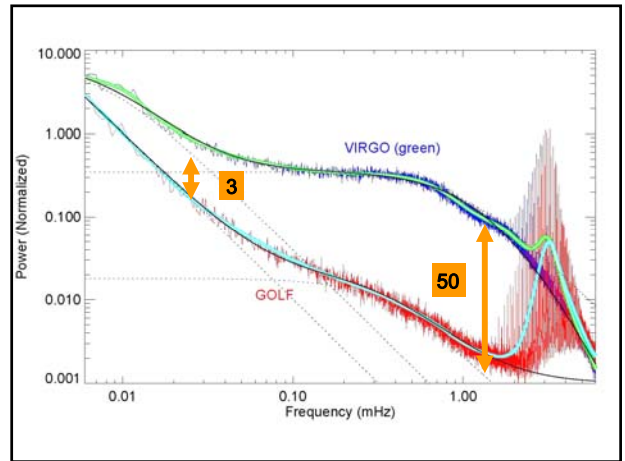
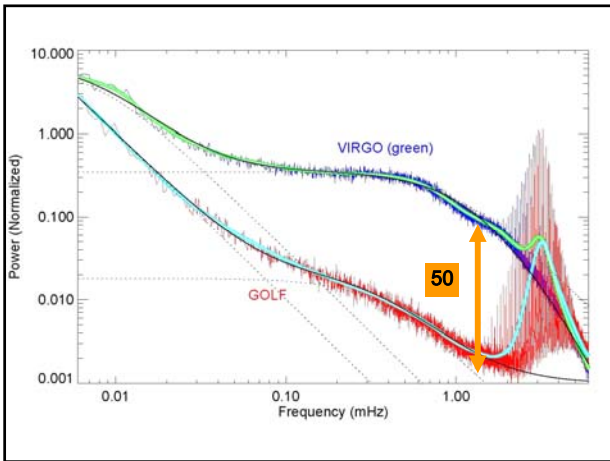
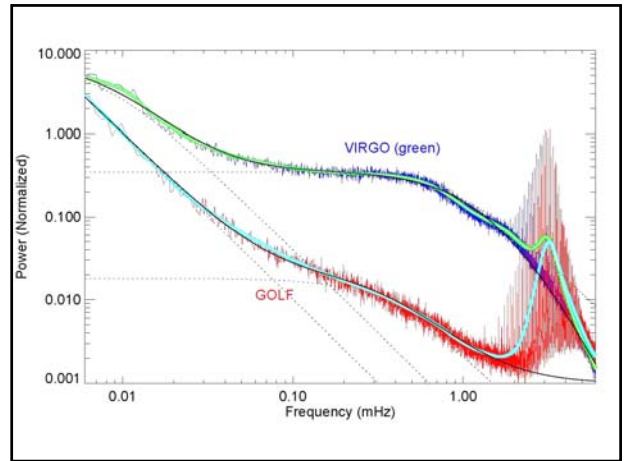
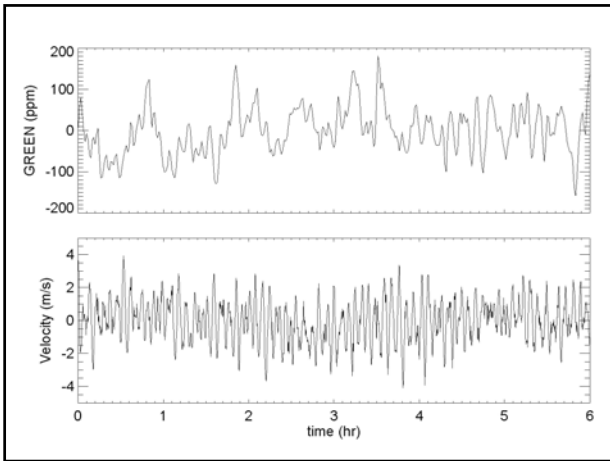
Damped and re-excited modes

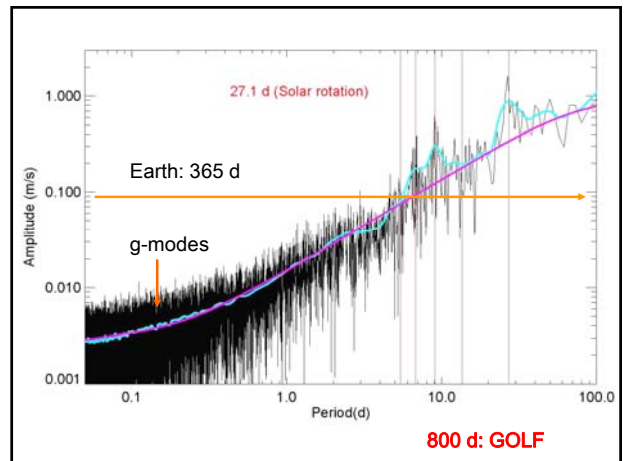
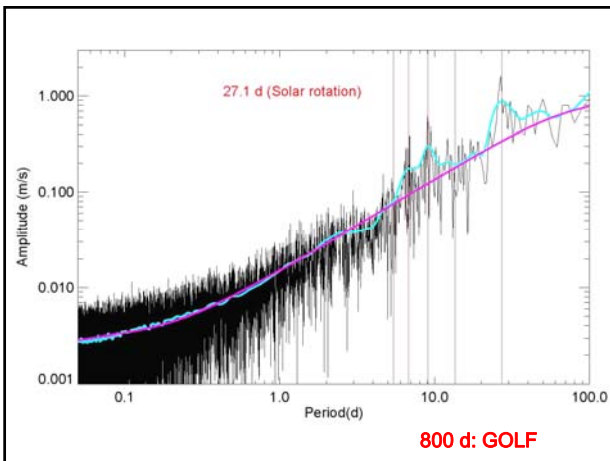
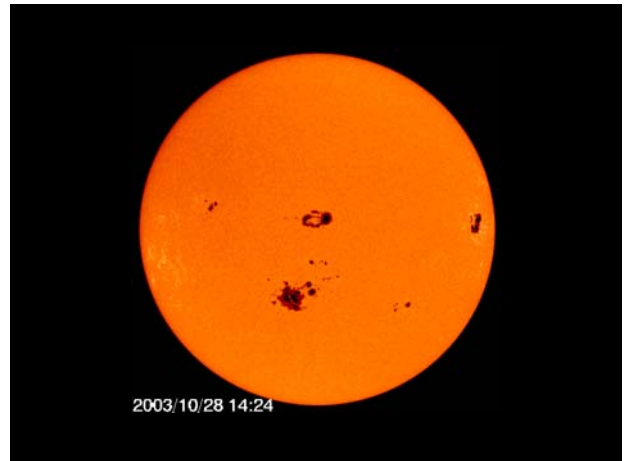
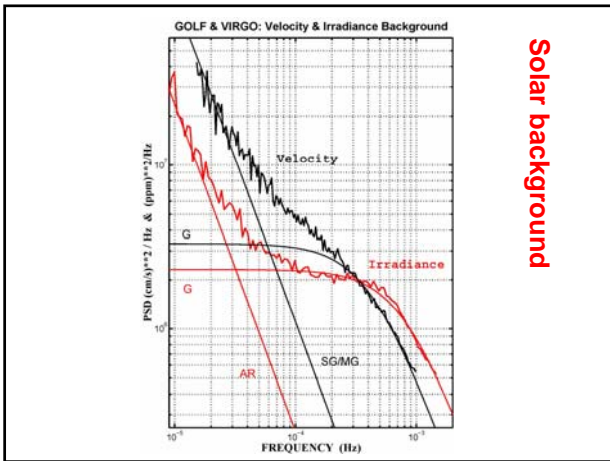
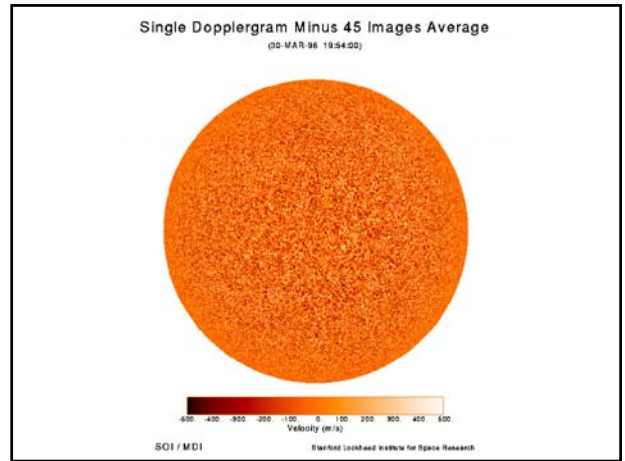
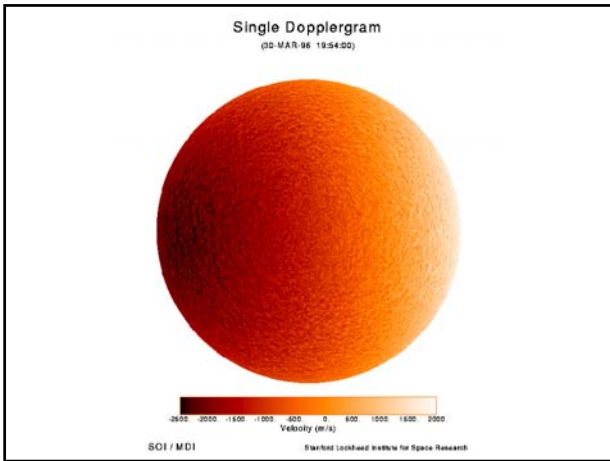
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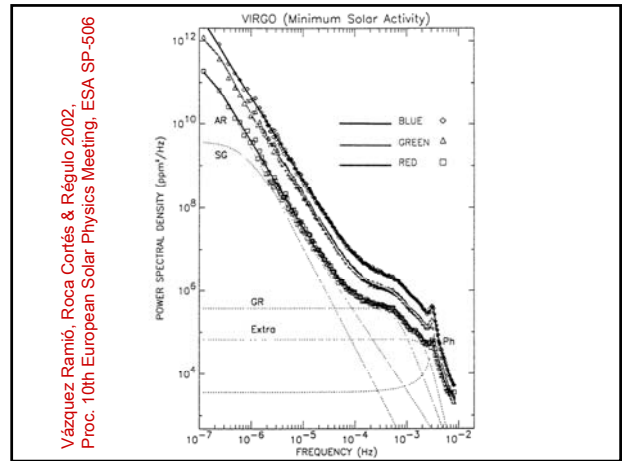
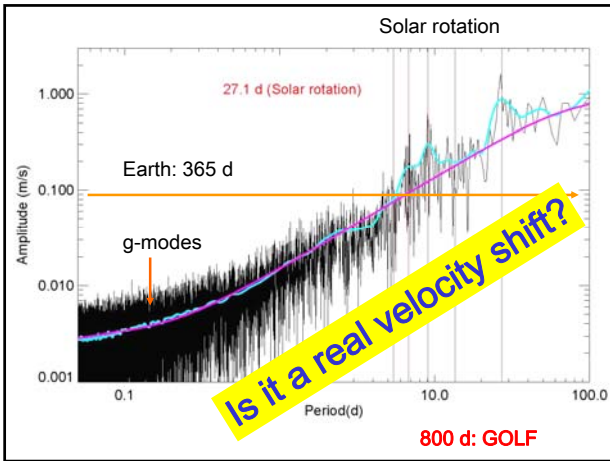




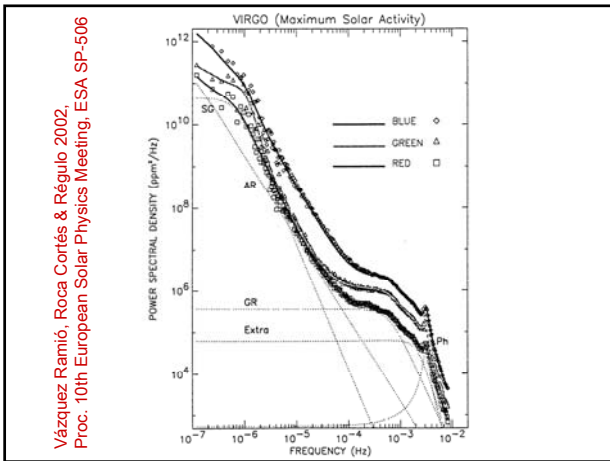




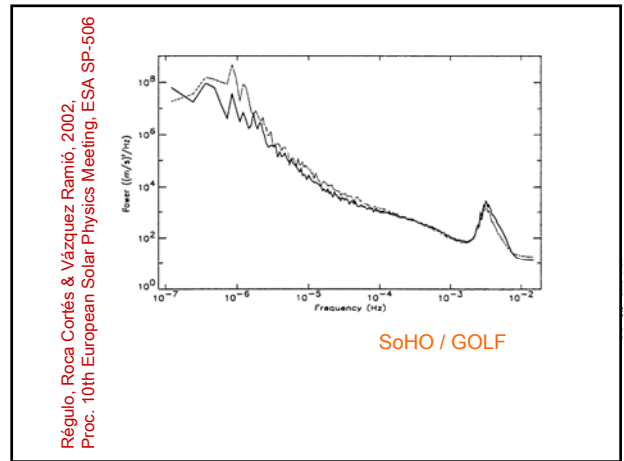




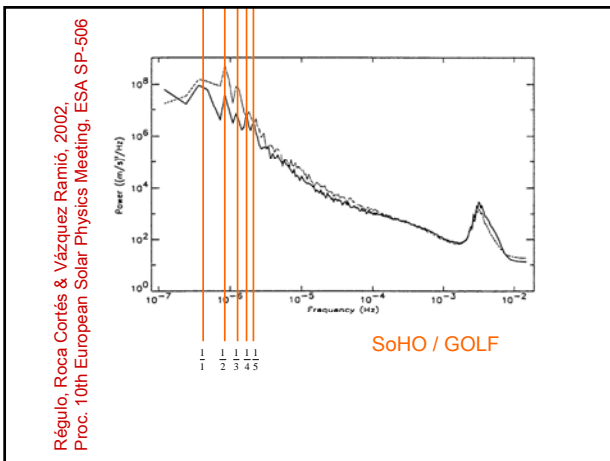
Vázquez Ramíó, Roca Cortés & Régulo 2002,
Proc. 10th European Solar Physics Meeting, ESA SP-506



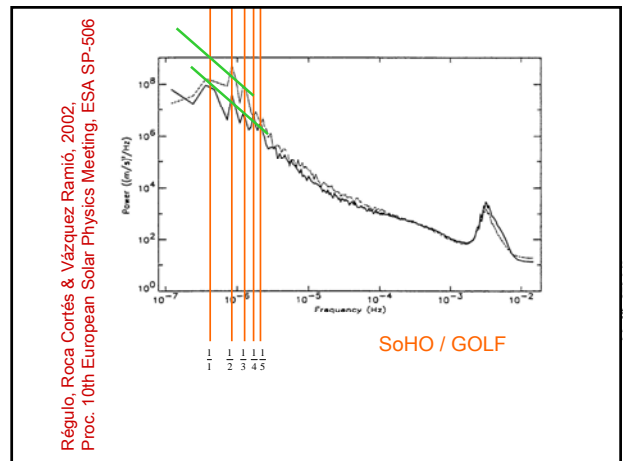
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